VSOP PROPOSAL COVER SHEETS

ID :

TR :

SR :

DEADLINE : 17 November, 1995
SEND TO : VSOP SOG, ISAS, 3-1-1 Yoshinodai, Sagamihara, Kanagawa 229, JAPAN

Please read Appendix C of Announcement of Opportunity for details on how to fill in this Cover Sheet.

(1) Date prepared : 17 Nov, 1995

(2) Proposal title : Distribution of Spectral Index in the CSS 3C380

(3)	INVESTIGATORS	INSTITUTION
P.I.	KAMENO Seiji	NAO, Japan
co-I.	INOUE Makoto	NAO, Japan
co-I.	TABARA Hiroto	Utsunomiya University, Japan
co-I.	FUJISAWA Kenta	ISAS, Japan
co-I.	WAJIMA Kiyoaki	ISAS, Japan
co-I.		
co-I.		
co-I.		
co-I.		

(4) Principal Investigator (or contact person) details...

Name : KAMENO Seiji
Address : 2-21-1 Osawa, Mitaka,
: Tokyo 181
: Japan
:

Internet :
kameno@hotaka.mtk.nao.ac.jp
Other e-mail :
Fax : +81-422-34-3869
Telephone : +81-422-34-3807

(5) Proposal Abstract :

Compact Steep-Spectrum sources (CSSs) have both intrinsically compact size (< 10 kpc) and steep spectrum ($\alpha \leq -0.5$). We found that CSSs also have active flat-spectrum core (Kameno et al. 1995). The steep-spectrum compact lobes dominate in integrated spectrum. Thus the spectrum index of jet emanating from flat-spectrum core rapidly get steeper toward lobes. To investigate how spectral steepening occures, we propose observations for the CSS 3C380 at 3 frequences and obtain distribution of spectral index in its jet. It consists of currently active flat-spectrum core and one-sided steep spectrum jet. A bright flat-spectrum knot 7 mas distant from the core is thought to be due to interaction between relativistic jets and local dense medium. Such a circumstance is preferable for our aim. Furthermore, a new knot component might be created at the flare of the core in 1993. We will also try to identify the new knot component.

(6) Proposal Category (indicate all that apply):

Object type:

☑ AGN, ☐ Masers, ☐ Stellar, ☐ Other :

Experiment type:

☑ Single-observation, ☐ Monitoring, ☐ Polarization,
☐ Time-critical, ☐ Target of Opportunity, ☐ Other :

(7) VSOP spacecraft observing mode (see Section 3 and Table 5 of the VSOP Proposer's Guide):

☑ 2 channel x 16 MHz, 2-bit (Standard mode),
☐ 2 channel x 32 MHz, 1-bit,
☐ 1 channel x 32 MHz, 2-bit

Phase calibration tones:

☑ On (Standard continuum mode),
☐ Off (Standard spectral line mode)

(Include justification of any non-standard choice at (14) below)

(8) Ground radio telescope setup

Polarization :

☑ VSOP Standard (IEEE LCP), ☐ Non-standard :

Recording mode :

☑ As for VSOP spacecraft (Standard), ☐ Other :

(9) Investigator participation in scheduling

☐ PI (or co-I) wishes to participate in scheduling ground radio telescopes
☐ PI (or co-I) wishes to participate in scheduling the space radio telescope

(10) Preferred correlator (see Sections 9.11 and 12 of VSOP Proposer's Guide):

☑ No preference, ☐ Mitaka, ☐ Socorro, ☐ Other :

(11) Preferred post-correlation data analysis location:

☐ Home Institution, ☑ Mitaka, ☐ NRAO AOC, ☐ JIVE, ☐ Other

(12) Post-correlation data analysis assistance required:

☑ None, ☐ Consultation, ☐ Extensive help

(13) Details of proposed experiments

An 'experiment' is one or more observations of one source in one wavelength band.
A request to observe the same source in all 3 wavelength bands requires 3 columns to be filled in.
To observe the same source at the same frequency multiple times – a 'monitoring experiment' – requires only one column to be filled in.
Number of experiments in this proposal: 3

	Experiment 1	Experiment 2	Experiment 3	Experiment 4
Source name	3C380	3C380	3C380	
RA (hh mm ss.s)	18 29 31.7	18 29 31.7	18 29 31.7	
Dec (dd mm ss)	+48 44 46	+48 44 46	+48 44 46	
J2000 or B1950?	J2000	J2000	J2000	
Observing frequency band (GHz)	1.6	5	22	
Continuum observations:				
Standard VSOP freq. channels?	☑	☑	☑	☐
Channel A range (MHz)				
Channel B range (MHz)				
Spectral line observations:				
Ch.A spectral line rest freq. (MHz)				
Ch.A LSR velocity (km/s)				
Ch.B spectral line rest freq. (MHz)				
Ch.B LSR velocity (km/s)				
Min. spectral channels per IF channel				
Correlator averaging time (sec)	2	2	2	
FWHM of field of view required (mas)	20	20	20	
No. of correlating passes (if >1)				
Measured total flux density (Jy)	14.7 (1.4 GHz)	7.50	3.18	
Measured correlated flux density on > 5000 km baseline (Jy)	0.40 (2.3 GHz)	1.01	0.5	
Image RMS needed (mJy/beam)	1.0	0.8	0.3	
Ground Radio Telescopes:				
Preferred choice:				
Number of medium telescopes	4	4	4	
Number of large telescopes	5	5	5	
Suggested array given at Item (14)	☐	☐	☐	☐
Minimum acceptable:				
Number of medium telescopes	2	2	2	
Number of large telescopes	2	2	2	
Suggested array given at Item (14)	☐	☐	☐	☐
Length of observation:				
Preferred length (orbits)	4	4	4	
Minimum acceptable length (orbits)	2	2	2	
Scheduling constraints:				
Preferred P.A. of beam *major* axis (deg)				
'No holes' (u,v) coverage?	☐	☐	☐	☐
Or maximum resolution (u,v) coverage?	☐	☐	☐	☐
Preferred range of dates for scheduling (for monitoring experiments give range for 1st observation only)	to	to	to	to
For monitoring programs:				
Number of observations				
Mean interval (days)				
Acceptable variance from mean (days)				

(14) Additional notes to the scheduler :

Please schedule all three experiments within a week.

(15) Attach a scientific and technical justification, not in excess of 2 pages of text and 2 pages of figures. Up to one page of (u,v) plots per source may optionally be included. (Refer to the VSOP Announcement of Opportunity for detailed instructions.) Preprints and reprints will not be forwarded to the Scientific Review Committee.

Send two paper copies of the complete proposal to:

VSOP Observing Proposals
VSOP Science Operations Group
Institute of Space and Astronautical Science
3-1-1 Yoshinodai, Sagamihara
Kanagawa 229 JAPAN

In addition, e-mail the completed LaTeX file to submit@vsopgw.isaslan1.isas.ac.jp

Cover Sheets of accepted proposals will be made available to the astronomical community.

Proposals must be received at ISAS by 17 November 1995